

Statistical Analysis of future Neutrino Mass Experiments including Neutrino-less Double Beta Decay

Werner Maneschg^{*} , Alexander Merle[†] , Werner Rodejohann[‡]

*Max-Planck-Institut für Kernphysik,
Postfach 10 39 80, 69029 Heidelberg, Germany*

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Abstract

We perform a statistical analysis with the prospective results of future experiments on neutrino-less double beta decay, direct searches for neutrino mass (KATRIN) and cosmological observations. Realistic errors are used and the nuclear matrix element uncertainty for neutrino-less double beta decay is also taken into account. Three benchmark scenarios are introduced, corresponding to quasi-degenerate, inverse hierarchical neutrinos, and an intermediate case. We investigate to what extend these scenarios can be reconstructed. Furthermore, we check the compatibility of the scenarios with the claimed evidence of neutrino-less double beta decay.

^{*}email: werner.maneschg@mpi-hd.mpg.de

[†]email: alexander.merle@mpi-hd.mpg.de

[‡]email: werner.rodejohann@mpi-hd.mpg.de

1 Introduction

Neutrino mass and lepton mixing represent an unambiguous proof that the Standard Model (SM) of elementary particles is incomplete. Various experiments with solar [1], atmospheric [2] and man-made [3, 4] neutrino sources imply non-trivial lepton mixing angles, as well as non-zero and non-degenerate neutrino masses. Their values are extremely suppressed with respect to the masses of the other (electrically charged) fermions of the SM. The most prominent and often studied mechanism to explain the smallness of neutrino masses is the see-saw mechanism [5]. The neutrino mass scale is here inversely proportional to the scale of its origin. In addition, lepton number violation is predicted: neutrinos are Majorana particles. Searching for this property will be a crucial test of the see-saw mechanism, but also of other mechanisms leading to small Majorana neutrino masses. Possible phenomenological consequences of lepton number violation are the generation of the baryon asymmetry of the Universe [6] or, at low energies, neutrino-less double beta decay ($0\nu\beta\beta$) [7]. This decay of certain nuclei, $(A, Z) \rightarrow (A, Z + 2) + 2 e^-$, which has not yet been observed, clearly violates lepton number by two units, and is intensively searched for [7]. We will assume here that light Majorana neutrinos are exchanged in the diagram responsible for $0\nu\beta\beta$. In this case, the amplitude for this process is proportional to the coherent sum

$$m_{ee} \equiv \sum_{i=1}^3 U_{ei}^2 m_i, \quad (1)$$

where m_i are the individual neutrino masses and U is the leptonic mixing, or Pontecorvo-Maki-Nakagawa-Sakata (PMNS), matrix. The absolute value of m_{ee} is called the effective mass. The entries U_{ei} can be written as $U_{e1} = \cos \theta_{12} \cos \theta_{13}$, $U_{e2} = \sin \theta_{12} \cos \theta_{13} e^{i\alpha}$ and $U_{e3} = \sin \theta_{13} e^{i\beta}$, where α and β are two currently unknown “Majorana phases” and $\theta_{12,13}$ are mixing angles. While θ_{13} is constrained mainly by short-baseline reactor experiments, θ_{12} is probed by solar and long-baseline reactor neutrino experiments. Their current best-fit values as well as 1σ and 3σ ranges can be obtained from three-flavor fits, the result being [8]

$$\sin^2 \theta_{12} = 0.32 (\pm 0.02)^{+0.08}_{-0.06}, \quad \sin^2 \theta_{13} = 0^{+0.019, 0.050}. \quad (2)$$

In what regards the neutrino masses, for a normal ordering one has $m_3 > m_2 > m_1$ with $m_2^2 = m_1^2 + \Delta m_{\odot}^2$ and $m_3^2 = m_1^2 + \Delta m_{\text{A}}^2$. In case of an inverted ordering we have $m_2 > m_1 > m_3$ with $m_2^2 = m_3^2 + \Delta m_{\odot}^2 + \Delta m_{\text{A}}^2$ and $m_1^2 = m_3^2 + \Delta m_{\text{A}}^2$. Here Δm_{\odot}^2 and Δm_{A}^2 are mass-squared differences with best-fit values and 3σ ranges $(7.9^{+1.1}_{-0.9}) \cdot 10^{-5} \text{ eV}^2$ and $(2.6^{+0.6}_{-0.6}) \cdot 10^{-3} \text{ eV}^2$, respectively [8]. Quasi-degenerate neutrino masses occur when $m_{1,2,3}^2 \gg \Delta m_{\text{A}}^2, \Delta m_{\odot}^2$. If neutrinos are Majorana particles, all low energy neutrino phenomenology can be described by the neutrino mass matrix $m_{\nu} = U^* m_{\nu}^{\text{diag}} U^{\dagger}$. It contains nine physical parameters. Seven out of the nine parameters of the neutrino mass matrix appear in $|m_{ee}|$. Therefore, it contains a large amount of information, in particular if complementary measurements of some of the other parameters exist. We also note that all parameters of m_{ν} , which do *not* influence neutrino oscillations show up in the effective mass. Those

are the the Majorana phases and, in particular, the individual neutrino masses (neutrino oscillations are only sensitive to mass-squared differences). For a review on the dependence of $|m_{ee}|$ on the various neutrino parameters see refs. [7, 9, 10] and references therein. In the present paper, in contrast to other works statistically analyzing future neutrino mass measurements including $0\nu\beta\beta$ [14, 11, 12, 13, 15, 16, 17], we focus on the neutrino mass scale, i.e. the value of the smallest neutrino mass. To this end we define three natural benchmark scenarios and investigate how future experiments may be able to constrain them. Our goal here is to combine as much mass-related information as possible.

2 Observables related to neutrino mass

Currently the strongest experimental limits¹ on the half-life of neutrino-less double beta decay are (all at 90 % C.L.) $1.9 \cdot 10^{25}$ y for ^{76}Ge [19] (see also [20]), $T_{1/2} \geq 3.0 \cdot 10^{24}$ y for ^{130}Te [21], $T_{1/2} \geq 5.8 \cdot 10^{23}$ y for ^{100}Mo and $T_{1/2} \geq 2.1 \cdot 10^{23}$ y for ^{82}Se [22]. The existing limits on $T_{1/2}$ will be improved considerably (by two orders of magnitude or more) in the near future by various experiments [7]. The uncertainty in nuclear matrix element (NME) calculations is a serious problem to translate these bounds into upper limits on the effective mass [23, 17]. We will take into account in particular this uncertainty in our analysis. Depending on the nuclei and NME, the current limit on the effective mass as extracted from the half-lifes given above lies between several tenths of and a few eV. This has to be compared with the predictions which can be made for the effective mass. Inserting the known ranges of the oscillation parameters, and varying the unknown parameters within their allowed ranges, one can generate plots as the ones in fig. 1.

They display (for $U_{e3} = 0$) the effective mass as a function of the smallest neutrino mass, the sum of neutrino masses

$$\Sigma \equiv \sum_{i=1}^3 m_i \quad (3)$$

and the kinematic neutrino mass

$$m_\beta \equiv \sqrt{\sum_{i=1}^3 |U_{ei}|^2 m_i^2}. \quad (4)$$

The latter two quantities can be measured through cosmological observations [24] and experiments like KATRIN [25], respectively. The latter experiment has a 5σ discovery potential of 0.35 eV for m_β , and a null result will lead to a 90 % C.L. limit of 0.2 or 0.17 eV [26]. In the sensitivity range of KATRIN, the relation $3m_\beta = \Sigma$ holds to a very good precision. Cosmology is expected to probe values of Σ down to the 0.1 eV range [24] (to be specific, we take a value of 0.15 eV in fig. 1). To achieve such impressive results, one takes

¹We note that there is a claimed positive signal for $0\nu\beta\beta$ from ref. [18]. We will turn to this issue later on.

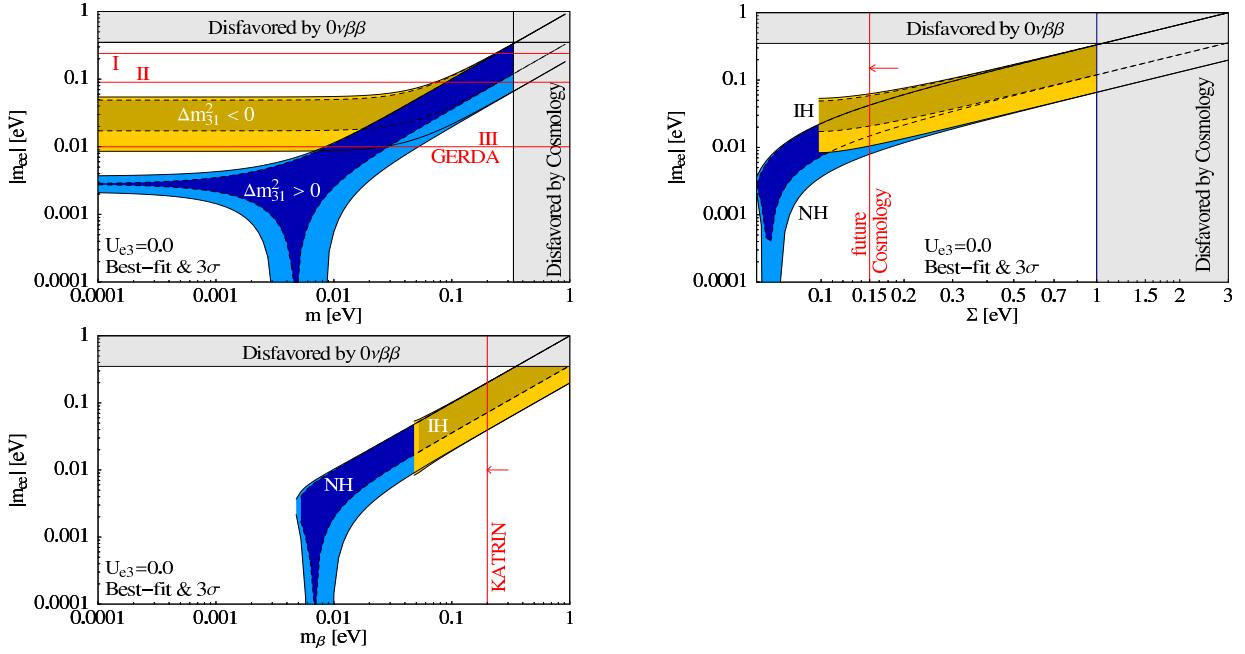


Figure 1: The effective mass as a function of the smallest neutrino mass, the sum of neutrino masses Σ and the kinematic neutrino mass m_β . The value $U_{e3} = 0$ and the current 3σ ranges of the other oscillation parameters have been used.

advantage of future observations of weak gravitational lensing of galaxies, and the cosmic microwave background or detailed analyses of the 21 cm hydrogen emission lines at high redshift. It is fair to say that a conservative limit on Σ is 1 eV. This value corresponds roughly to the bound obtained from WMAP 5-year data alone [28]. Recall that neutrino mass bounds from cosmology depend strongly on the data sets, the priors and the model, i.e., adding parameters which are degenerate with neutrino masses will relax the bounds, see, e.g., [29]. Finally, current limits for m_β are 2.3 eV [27].

The blue and yellow bands in fig. 1 correspond to the normal and inverted mass ordering of the neutrinos, respectively. The darker areas in the blue and yellow bands are obtained when the oscillation parameters are fixed to their best-fit values and only the Majorana phases are varied. The lighter areas correspond to the 3σ ranges of the oscillation parameters. Note that this broadening is very weak for the maximum value of $|m_{ee}|$ in the case of inverted mass ordering and for quasi-degenerate neutrinos. This is because the upper limits on $|m_{ee}|$ are roughly $\sqrt{\Delta m_A^2}$ and m_3 , respectively, and varying the oscillation parameters has very little impact. In the first plot of fig. 1, we have indicated three special values of $|m_{ee}|$ which correspond to the goals of the three phases of the GERDA experiment (where a certain NME has been assumed, see [30] for details).

3 Statistical analysis

Now we will perform a statistical analysis to investigate how well it will be possible to reconstruct different realistic physical scenarios with upcoming neutrino mass experiments. Note that, since we want to investigate realistic situations, we concentrate only on cases that can be probed in the near future. For definiteness, we consider the inverted mass ordering and three different scenarios called $Q\mathcal{D}$ (quasi-degenerate), $I\mathcal{N}\mathcal{T}$ (intermediate) and $I\mathcal{H}$ (inverted hierarchy) that are defined by different values of the smallest neutrino mass m_3 . Note that the $Q\mathcal{D}$ scenario would, to very large extent, also apply to a normal mass ordering. The hypothetical “true values” for the different observables in these scenarios are:

Scenario	m_3 [eV]	$ m_{ee} $ [eV]	m_β [eV]	Σ [eV]
$Q\mathcal{D}$	0.3	$0.11 - 0.30$	0.30	0.91
$I\mathcal{N}\mathcal{T}$	0.1	$0.04 - 0.11$	(0.11)	0.32
$I\mathcal{H}$	0.003	$0.02 - 0.05$	(0.05)	(0.10)

We have used here the best-fit values for the oscillation parameters. The range for $|m_{ee}|$ originates from the variation of the Majorana phases α and β . Note that the KATRIN experiment will only be able to measure m_β in the case of the $Q\mathcal{D}$ scenario, while for $I\mathcal{N}\mathcal{T}$ and $I\mathcal{H}$ it will only provide an upper limit. The same is true for the measurement of Σ in the $I\mathcal{H}$ scenario. These cases are indicated in the table by writing the respective values in brackets.

Let us now give a summary of the different experimental errors and theoretical uncertainties. Regarding the error on the effective mass in $0\nu\beta\beta$, we have to distinguish between experimental and “theoretical uncertainties”, where the latter result from the NME uncertainty. The experimental error can be included by noting that the decay width depends quadratically on the effective mass. Thus,

$$\sigma(|m_{ee}|_{\text{exp}}) = \frac{|m_{ee}|_{\text{exp}}}{2} \frac{\sigma(\Gamma_{\text{obs}})}{\Gamma_{\text{obs}}}, \quad (5)$$

where $|m_{ee}|_{\text{exp}}$ is the measured value of the effective neutrino mass and $\sigma(\Gamma_{\text{obs}})$ is the experimental error on the measured decay width Γ_{obs} for neutrino-less double beta decay. For definiteness, we choose the ratio of the latter two as

$$\frac{\sigma(\Gamma_{\text{obs}})}{\Gamma_{\text{obs}}} \simeq 23.3\%, \quad (6)$$

which is the value obtainable in the GERDA experiment [30]. We combine, similarly to the procedure developed in ref. [11], the experimental error with the theoretical NME error via

$$\sigma(|m_{ee}|) = (1 + \zeta) \left(|m_{ee}| + \sigma(|m_{ee}|_{\text{exp}}) \right) - |m_{ee}|, \quad (7)$$

where $\zeta \geq 0$ parameterizes the NME uncertainty and $\sigma(|m_{ee}|_{\text{exp}})$ is given in eq. (5). Following ref. [14], we define a covariance matrix

$$S_{ab} \equiv \delta_{ab} \sigma^2(a) + \sum_i \frac{\partial T_a}{\partial x_i} \frac{\partial T_b}{\partial x_i} \sigma_i^2, \quad (8)$$

where $T_1 = |m_{ee}|$, $T_2 = \Sigma$ and $T_3 = m_\beta^2$. Furthermore, $\sigma^2(a)$ is the error on T_a , and a, b label the entries in the covariance matrix. The x_i are the oscillation parameters that enter $|m_{ee}|$ (and m_β , though in the observable range of m_β they have basically no influence). The errors on the T_a are given by eq. (7) as well as by $\sigma(m_\beta^2) = 0.025 \text{ eV}^2$ [25, 26] and $\sigma(\Sigma) = 0.05 \text{ eV}$ [24].

Defining $v_a = T_a - (T_a)_{\text{exp}}$, where $(T_a)_{\text{exp}}$ denotes the experimental value of T_a , our χ^2 -function to be minimized is

$$\chi^2 = v^T S^{-1} v. \quad (9)$$

All oscillation parameters are set to their current best-fit values and their (symmetrized) standard deviations are determined from their 1σ -ranges, which is a good approximation for future 3σ -ranges. Anyway, the impact of different numerical values here would not lead to qualitatively different results. We first minimize the χ^2 from eq. (9) with respect to the Majorana phases α and β . The resulting function is $\chi_{\text{res}}^2 = \min_{\alpha, \beta} \chi^2$. We then continue by plotting the resulting 1σ , 2σ and 3σ ranges for the smallest neutrino mass m_3 determined by setting $\Delta\chi^2 = \chi_{\text{res}}^2 - \chi_{\text{res},\text{min}}^2$ equal to 1, 4 and 9. This corresponds to a χ^2 -function with one free parameter (namely m_3). $|m_{ee}|_{\text{exp}}$ is the assumed measured value of $|m_{ee}|$, on which the reconstructed range of m_3 depends. The minimum in the $|m_{ee}|_{\text{exp}}-m_3$ plane is determined such that $\Delta\chi^2$ is zero in the true region of the corresponding scenario (e.g., \mathcal{QD}).

The results of our analysis are shown as the solid lines in the left column of figs. 2, 3 and 4. In all cases, we have calculated the result for a consistent measurement (i.e., m_β and Σ are measured at their true values in the corresponding scenarios). The NME uncertainties we have chosen are $\zeta = 0$ (no uncertainty), 0.25 and 0.5. We have checked that values of $\zeta > 0.5$ will lead to results not too much different from the ones for $\zeta = 0.5$. The value $\zeta = 0.25$ is a quite typical one, cf. refs. [23, 17]. This uncertainty arises from the highly non-trivial calculations of the nuclear part of the neutrino-less double beta decay process. Different methods, and even different Ansätze within the same framework, differ in their result, and their spread is commonly taken into account as “theoretical uncertainty”. Glancing at Fig. 5 in ref. [31], where the results of different methods of the NME calculation are compared for different nuclei including ^{76}Ge , one can indeed see that the spread of the respective values around their mean value is about 0.2. We conclude that the values we use are realistic and typical.

The true values of $|m_{ee}|$ and m_3 are marked by the vertical black lines. The plots illustrate how well we can reconstruct the different scenarios for the various values of NME uncertainty. Having a look at fig. 2, we see that the \mathcal{QD} scenario can be reconstructed quite well, which is not surprising since in that case the KATRIN experiment as well as the cosmological measurement will provide a non-trivial signal. E.g., for $|m_{ee}|_{\text{exp}} = 0.20 \text{ eV}$,

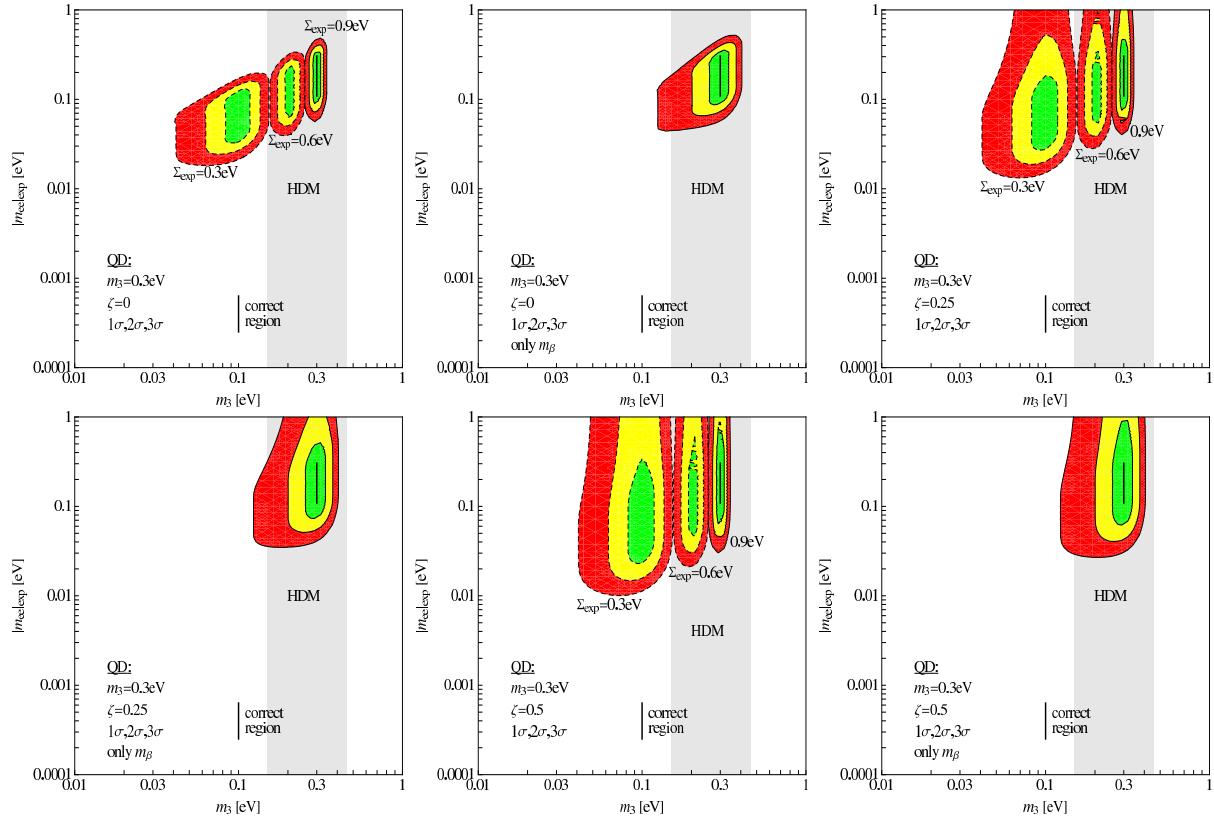


Figure 2: 1σ , 2σ and 3σ regions in the m_3 - $|m_{ee}|_{\text{exp}}$ plane for the QD scenario. The left column shows the correct (solid line) as well as two possible incorrect cosmological measurements (dashed lines). The less desirable case, namely only taking into account a KATRIN measurement, is shown in the plots on the right. The area denoted HDM is the range of $|m_{ee}|$ from the claim of part of the Heidelberg-Moscow collaboration.

the 1 , 2 and 3σ ranges for m_3 are $0.28 - 0.32$ eV, $0.27 - 0.33$ eV and $0.25 - 0.35$ eV, while the true value is 0.30 eV. Therefore, the reconstruction is quite accurate. This remains true also if the uncertainty in NME is non-zero because the plots are still narrow around the true value of m_3 (the numerical values suffer nearly no change) even though, with a larger NME uncertainty, also higher values of $|m_{ee}|_{\text{exp}}$ are plausible. This is true for all three scenarios under consideration.

Similar statements hold for the INT scenario shown in fig. 3, even though m_β cannot be measured now. However, because there will still be a measurement of Σ , we have sufficient information on the neutrino mass. In case the central measured value is $|m_{ee}|_{\text{exp}} = 0.08$ eV and $\zeta = 0$ the ranges are $0.08 - 0.12$ eV at 1σ and $0.05 - 0.15$ eV at 3σ . In case of $\zeta = 0.5$ we find $0.08 - 0.12$ eV at 1σ and $0.04 - 0.15$ eV at 3σ . The mass scale has now a 3σ uncertainty of 50 %, to be compared with roughly 15 % in the QD scenario.

For IH , in turn, there is no measurement that gives information m_3 . Hence, it is only

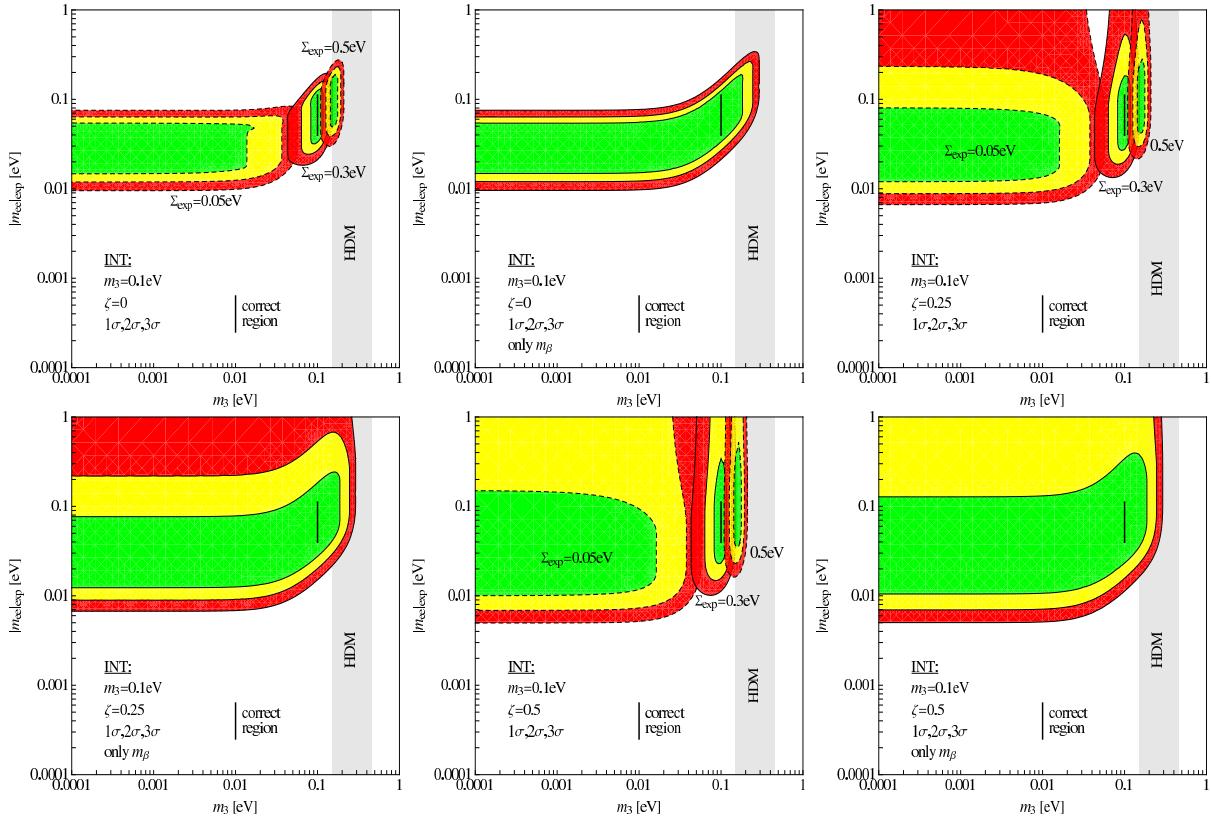


Figure 3: Same as fig. 2 for the \mathcal{INT} scenario.

possible to give an upper limit on the smallest neutrino mass, as illustrated by the long horizontal band in the left column of fig. 4. Note that this band corresponds to the yellow band marking the inverted mass ordering in the upper plot of fig. 1. This upper limit is almost trivial, i.e., it corresponds to the neutrino mass limit obtainable from $0\nu\beta\beta$ alone. To give some numerical values, for $|m_{ee}|_{\text{exp}} = 0.04$ eV one would have the 1 (3) σ ranges $m_3 < 0.03$ (0.07) eV for $\zeta = 0$ and for $\zeta = 0.5$. Due to the bound on Σ , there is very little dependence on ζ .

Up to now, the discussion has focused on the case in which all measurements are compatible. As an example for inconsistency we discuss here a possible clash between results from KATRIN and from cosmology. To this end we leave $(m_\beta)_{\text{exp}}$ equal to the true value of the corresponding scenario (new physics is not expected to influence m_β [32]) and take values of Σ_{exp} which are smaller or larger than the true value. There are many scenarios or models in the literature which can lead to wrong values of Σ , see, e.g., refs. [33]. The result is shown by the areas within the dashed lines in the left columns of figs. 2-4. Having a look at $Q\mathcal{D}$ first, we realize immediately that the physical range is reconstructed incorrectly. Hence, if there are systematic errors in the cosmological measurement, or unknown features

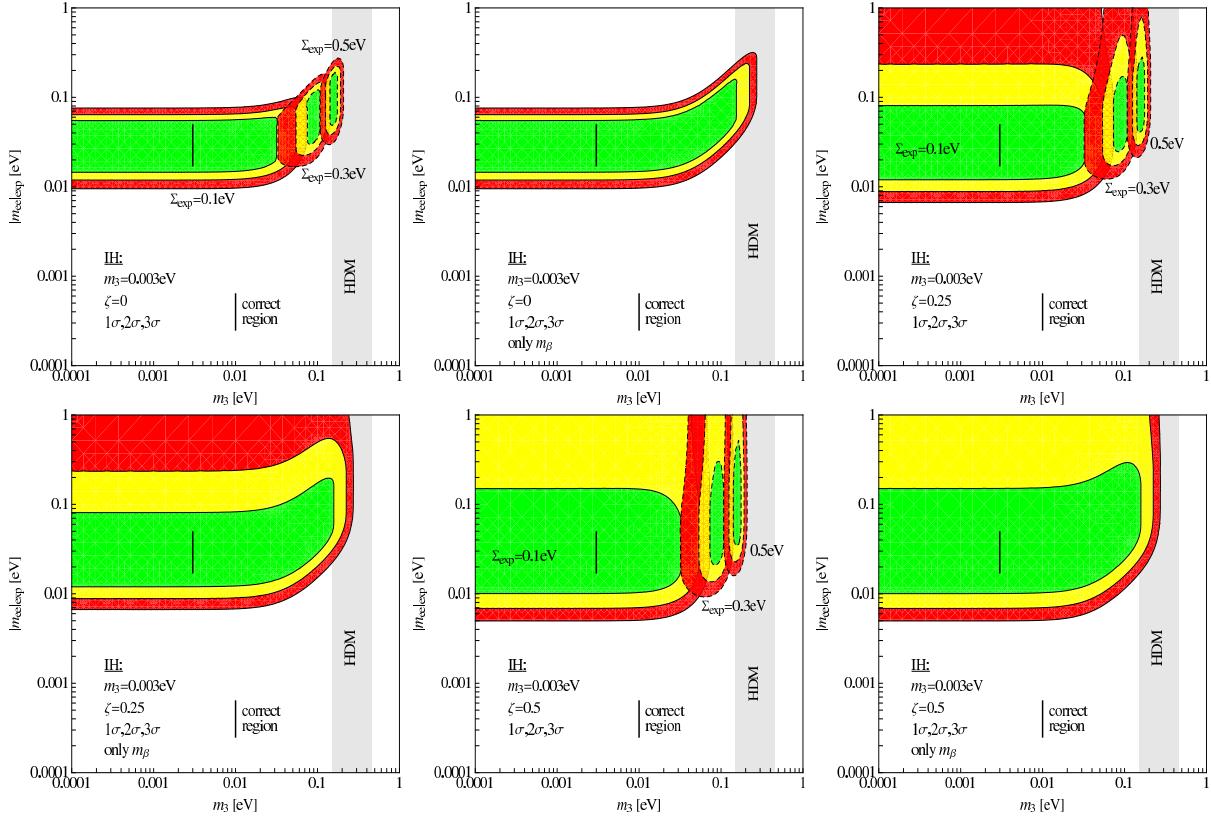


Figure 4: Same as fig. 2 for the \mathcal{IH} scenario.

in cosmology which we are not aware of, a wrong neutrino mass is reconstructed. In the \mathcal{QD} case there is still information from KATRIN, which leads to a reconstructed neutrino mass at most one order away from the true value, even if the wrong Σ is taken into account. For the \mathcal{INT} scenario, however, there is no information from KATRIN. Consequently, it might be that a *wrong* upper limit on m_3 is concluded, as illustrated by the long band for $\Sigma_{\text{exp}} = 0.05$ eV in the upper left plot of fig. 3. This is an example wherein one could draw a wrong conclusion by taking the cosmological measurement at face value. As expected, even worse cases may exist for the \mathcal{IH} scenario. E.g., in the upper left plot of fig. 4 one would, for $\Sigma_{\text{exp}} = 0.3$ eV, reconstruct a smallest neutrino mass of roughly 0.1 eV, to be compared with the true value of $m_3 = 0.003$ eV. For the \mathcal{IH} scenario, one might not even realize that there is an inconsistency, since in that case, the KATRIN experiment can only provide an upper limit which is too far away from the true value of m_3 .

One possible cross-check (or the possible consequence if one indeed finds that the results from KATRIN and from cosmology do not fit together) would be to dismiss the cosmological data altogether. We have also analyzed this case. Here, S_{ab} from eq. (8) as well as v_a would change from 3-dimensional to 2-dimensional objects while the rest of the procedure remains the same. The results for this analysis are plotted in the right columns of figs. 2-4,

again for different values of the NME uncertainty. For \mathcal{QD} , the most optimal scenario, neglecting cosmology, would simply increase the errors in the determination of m_3 : e.g., for $|m_{ee}|_{\text{exp}} = 0.20$ eV and $\zeta = 0$ the ranges are $0.26 - 0.34$ eV at 1σ and $0.16 - 0.41$ eV at 3σ , while for $\zeta = 0.5$ we find $0.26 - 0.34$ eV at 1σ and $0.13 - 0.41$ eV at 3σ . The NME uncertainty has now a slightly bigger impact, and the error on m_3 increases by a factor of three, since now it is about 50 % while it was roughly 15 % when Σ has been included in the analysis. For the \mathcal{INT} scenario, however, there is a major difference to the former case: since now there is no other measurement besides $|m_{ee}|_{\text{exp}}$ providing information on m_3 , we can only derive an upper limit instead of determining a certain range for m_3 . This is indicated by the band in the upper right plot of fig. 3. Finally, for \mathcal{IH} , the limit on m_3 gets only slightly worse compared to the case of a Σ , which is too small to be measured. In this case there would not even be a real drawback in taking into account the KATRIN result only. It remains to be said that in all cases a higher uncertainty for the NME does not significantly modify the conclusions in what concerns the value of m_3 . Finally, it is worth mentioning that if in \mathcal{QD} scenarios the error on Σ is decreased (increased), the obtained error on the neutrino mass is decreased (increased) by approximately the same factor.

With our analysis we can also compare the compatibility of our three benchmark scenarios with the range for m_3 of $0.15 - 0.46$ eV, calculated as the (global fit) 2σ range in ref. [16] from the claim in ref. [18]. We give the implied range for m_3 as the gray band in figs. 2, 3 and 4. We see that scenario \mathcal{QD} is consistent with the claim, even for a measurement of $\Sigma = 0.6$ eV, to be compared with the true value $\Sigma = 0.9$ eV. The \mathcal{INT} scenario (\mathcal{IH} scenario) is barely (very) incompatible for measured “true” values, but a too high value of Σ_{exp} can lead again to compatibility. We see that testing the claim and comparing it with cosmology is a non-trivial task (see also [17]).

4 Conclusions

In this work we have investigated possible constraints on the neutrino mass in future experiments. We assumed realistic errors on the observables, in particular for neutrino-less double beta decay. Then, we have checked how certain realistic benchmark scenarios, which correspond to different regimes for the smallest neutrino mass, can be reconstructed from future measurements. Furthermore, we have pointed out how wrong conclusions could be drawn from inconsistent results, i.e., if cosmology provides a wrong value for the sum of neutrino masses. In case of consistent measurements we may summarize as follows: typical 3σ errors for quasi-degenerate neutrino masses range from roughly 15 % (including Σ) to 50 % (excluding Σ), where NME uncertainties play a larger role in the latter case. Intermediate scale masses can also be determined with 50 % uncertainty. In case of an inverted hierarchy, the effective mass is constant for a large range of the smallest mass, which allows only to derive upper limits on it.

Acknowledgments

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